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THE CRITICAL MASS RATIO FOR W UMA-TYPE CONTACT BINARY SYSTEMS (Serb. Astron. J. № 208 (2024), 1)

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In the paper by Arbutina and Wadhwa (2024), due to an error at the production stage, axes designations are incorrect or missing in Figs. 2–5. The corrected Figures are reproduced below.

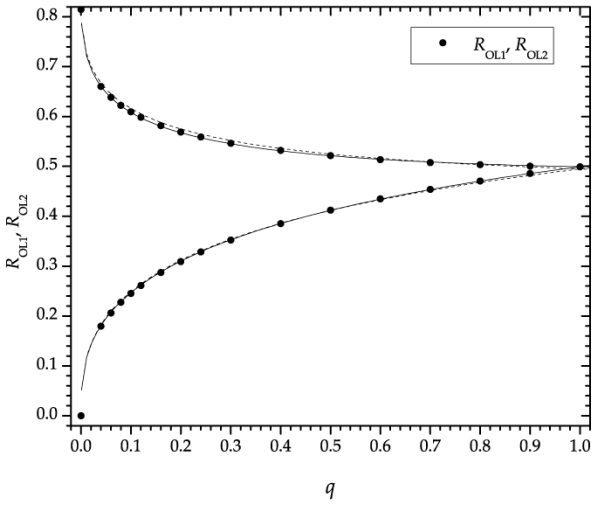

Fig. 2: Mean radii for outer Roche lobes $R_{\rm OL1}$ and $R_{\rm OL2}$. Filled circles are numerical data from Mochnacki (1984) tables, dashed curves are approximations given by Yakut and Eggleton (2005), while solid curves represent our approximations from Eq. (7).

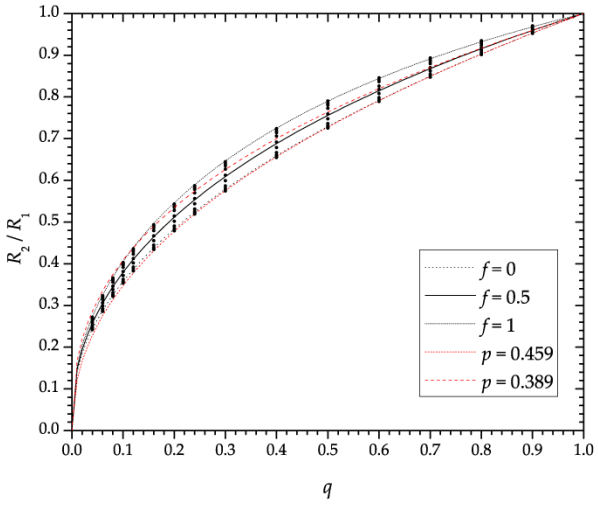

Fig. 3: The numerical data for the ratio R_2/R_1 from Yakut and Eggleton (2005) and our approximations. The dotted curve represents a power-law fits $R_2/R_1 \approx q^p$ to the observational data ($p = 0.459$, Kuiper 1941, Csizmadia and Klagyivik 2004) ($p = 0.389$, Poro et al. 2022).

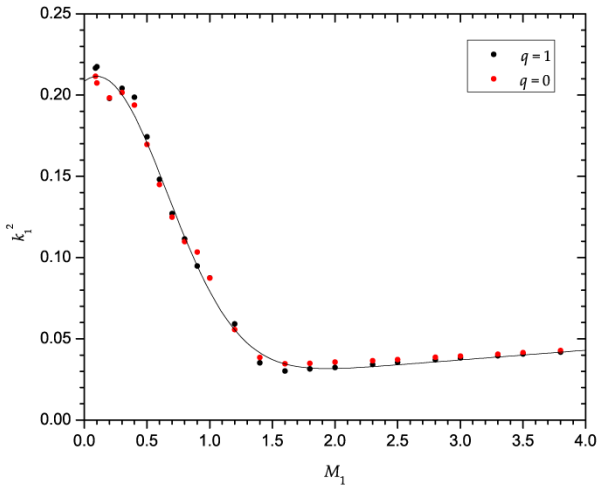

Fig. 4: The data (Landin et al. 2009) and $k_1^2 = k_1^2(M_1)$ relation fit.

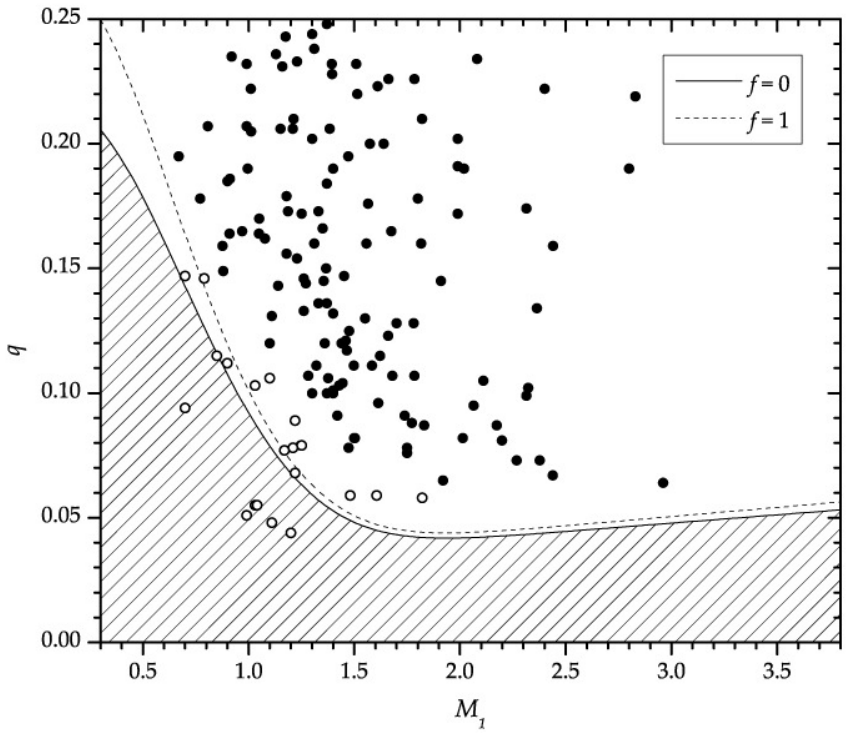

Fig. 5: The instability mass ratio versus primary mass. The data are from Latković et al. (2021) (filled circles) and Table 2 (open circles).

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